

Section 7.4 -7
Linear theory of waves & instability

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Basic Seminar IIA

7.4 Perturbations at a two-fluid interface

Irrotational Flow of **barotropic ideal** fluid with **conservative force**

- "Irrotational" ... $\underline{\Gamma \equiv \nabla \times \mathbf{v} = 0}$
 - \mathbf{v} can be written as the gradient of a certain scalar potential
 $\underline{\mathbf{v} = \nabla \phi}$ ← we can use ϕ instead of \mathbf{v}

Kelvin's vorticity theorem

- for barotropic ideal fluid with conservative body force

$$D\Gamma/Dt = 0$$

- **Euler eq. (Navier-Stokes eq. for ideal(=inviscid) fluid)** under conservative force

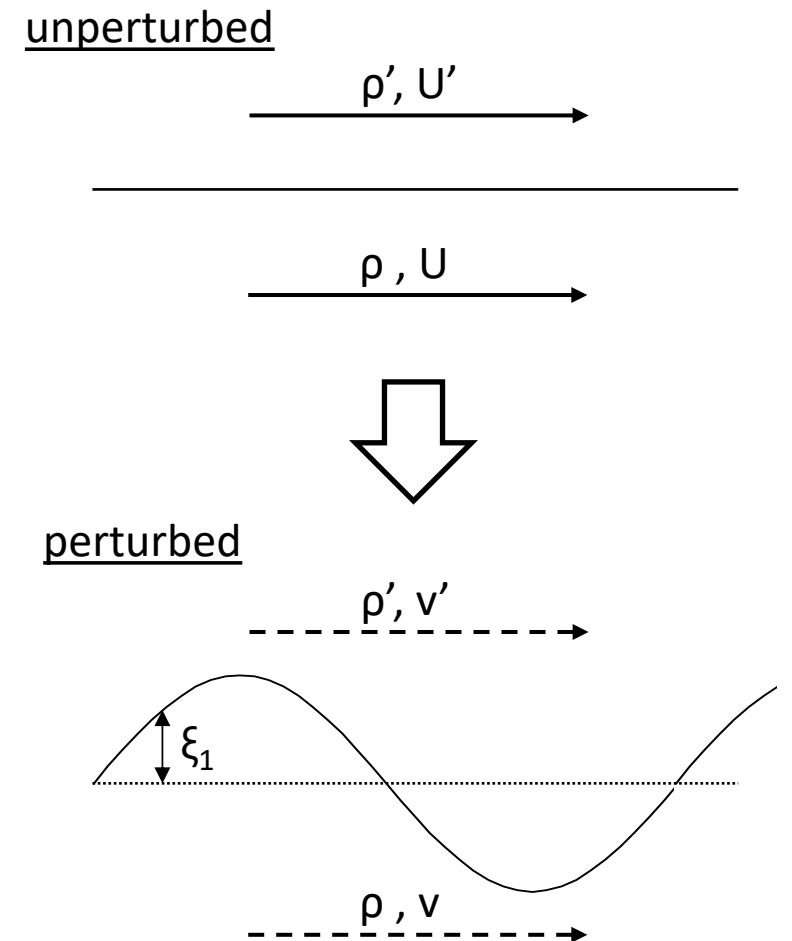
$$-\frac{\partial \phi}{\partial t} + \frac{1}{2}v^2 + \frac{p}{\rho} + \Phi = \boxed{F(t)} \quad \dots (7.29)$$

does not depend on position

7.4 Perturbations at a two-fluid interface

• Assumptions

- **Incompressible** ($\Leftrightarrow \nabla \cdot \mathbf{v} = 0 \Leftrightarrow$ continuity eq.)
 - each fluids have constant densities:
 ρ and ρ' (which do **not vary with position.**)
- **Irrotational** ($\Gamma \equiv \nabla \times \mathbf{v} = 0$)
 - $\mathbf{v} = -\nabla \phi$
- **Conservative force** ($\mathbf{F} = \nabla \Phi$)
 - Note: Be sure to distinguish between ϕ and Φ !!!



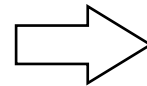
7.4 Perturbations at a two-fluid interface

- **formulation**

- general config.

velocity in lower/upper fluid:

$$\begin{aligned} v(x,z,t) &= \underbrace{U}_{\text{unperturbed}} + \underbrace{v_1}_{\text{perturbed}}, \\ v'(x,z,t) &= \underbrace{U'}_{\text{unperturbed}} + \underbrace{v'_1}_{\text{perturbed}}. \end{aligned}$$

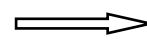


velocity potentials:

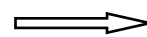
$$\begin{aligned} \phi &= \underbrace{-U x}_{\text{unperturbed}} + \underbrace{\phi_1}_{\text{perturbed}}, \\ \phi' &= \underbrace{-U' x}_{\text{unperturbed}} + \underbrace{\phi'_1}_{\text{perturbed}}. \end{aligned}$$

- boundary conditions

1. the perturbations will vanish @ far away from the interface
2. Pressure has to be continuous across the interface



$$\phi_1, \phi'_1 \rightarrow 0 \quad (@ \quad z \rightarrow \infty / -\infty)$$



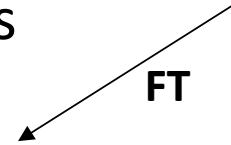
$$P(z=0) = P'(z=0)$$

7.4 Perturbations at a two-fluid interface

• formulation

$$\phi'_1 \rightarrow 0$$

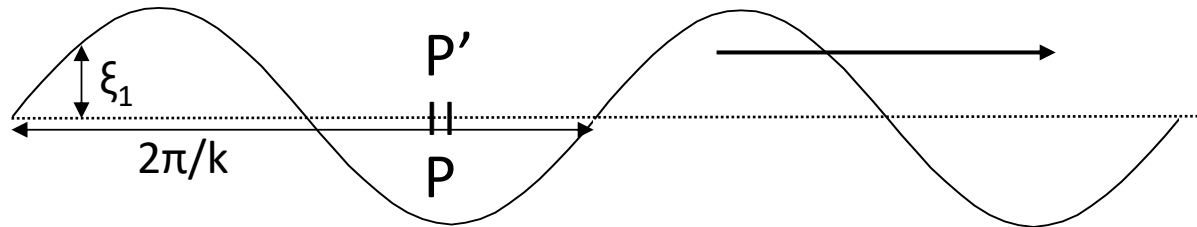
periodic symmetry $\rightarrow \xi_1 = f(t) \cdot \exp(ikx)$
 along x-axis



$$\xi_1(\omega) = A \exp(-i\omega t) \cdot \exp(ikx)$$

... (7.36)

ρ', ϕ'



ρ, ϕ

$$\phi_1 \rightarrow 0$$

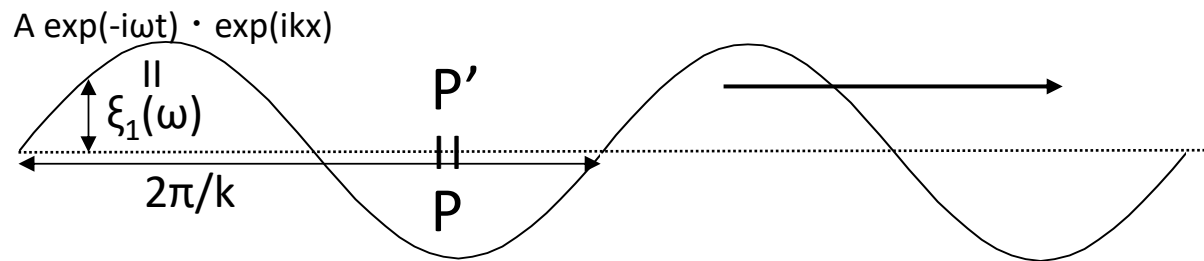
if we can linearize the perturbation eq.,
each mode evolve separately.
 Then we can consider one mode
 as a representative.

7.4 Perturbations at a two-fluid interface

• formulation

$$\phi'_1 \rightarrow 0$$

ρ', ϕ'



ρ, ϕ

$$\phi_1 \rightarrow 0$$

Linearized Perturbation eq.

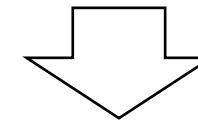
vertical velocity at vicinity of $z=0$

$$-\frac{\partial \phi_1}{\partial z} = \frac{\partial \xi_1}{\partial t} + U \frac{\partial \xi_1}{\partial x} \quad \dots (7.34,35)$$

Poisson eq.

Incompressible & Irrotational

$$\rightarrow \nabla^2 \phi = 0 \quad \dots (7.31)$$



$$\phi_1 = C \exp(-i\omega t) \cdot \exp(ikx + kz) \quad \dots (7.37)$$

$$\phi'_1 = C' \exp(-i\omega t) \cdot \exp(ikx - kz) \quad \dots (7.38)$$

7.4 Perturbations at a two-fluid interface

- **Boundary condition at the interface**

from (7.29) : $-\frac{\partial\phi}{\partial t} + \frac{1}{2}v^2 + \frac{p}{\rho} + \Phi = F(t)$

we can calculate pressure **P** & **P'** at z=0

and obtain (7.45) : $\rho \left(-\frac{\partial\phi_1}{\partial t} - U\frac{\partial\phi_1}{\partial x} + g\xi_1 \right) = \rho' \left(-\frac{\partial\phi'_1}{\partial t} - U'\frac{\partial\phi'_1}{\partial x} + g\xi_1 \right)$ at vicinity of z=0
... (7.45)

substituting ξ_1 , ϕ_1 , ϕ'_1 into

ξ - ϕ relation eq. (= 7.34, 35) and this boundary condition (7.45),
finally we get following **dispersion relation**.

$$\frac{\omega}{k} = \frac{\rho U + \rho' U'}{\rho + \rho'} \pm \left(\frac{g}{k} \frac{\rho - \rho'}{\rho + \rho'} - \frac{\rho \rho' (U - U')^2}{(\rho + \rho')^2} \right)^{1/2} \quad \dots (7.48)$$

7.4 Perturbations at a two-fluid interface

- **Summary of Linear Analysis for perturbations at a two-fluid interface**

- perturbed interface

$$\dots \xi_1 = A \exp(-i\omega t) \cdot \exp(ikx)$$

- dispersion relation

$$\dots \frac{\omega}{k} = \frac{\rho U + \rho' U'}{\rho + \rho'} \pm \left(\frac{g \rho - \rho'}{k \rho + \rho'} - \frac{\rho \rho' (U - U')^2}{(\rho + \rho')^2} \right)^{1/2} \equiv D$$

$D > 0 \rightarrow$ perturbations will not grow

ex.) $\rho > \rho'$, $U = U' = 0$... **surface gravity wave**

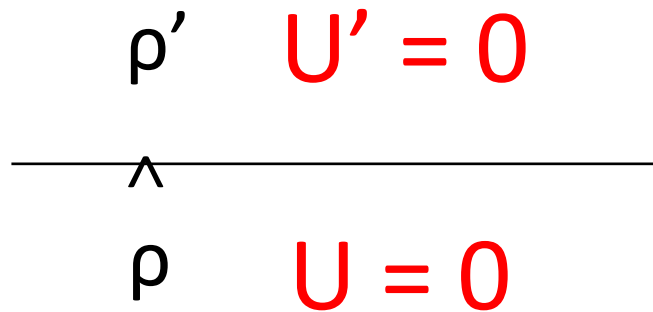
$D < 0 \rightarrow$ perturbations will grow

ex1.) $\rho < \rho'$... **Rayleigh-Taylor instability**

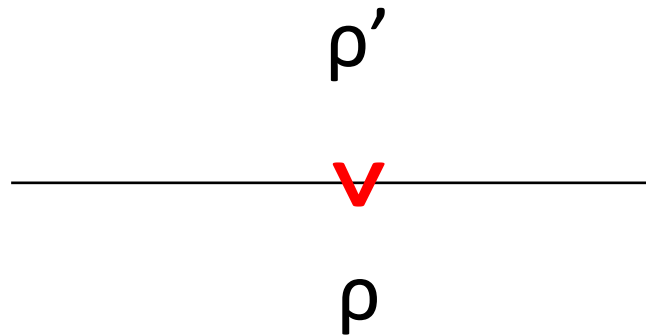
ex2.) $\rho > \rho'$... **Kelvin-Helmholtz instability**

7.4 Perturbations at a two-fluid interface

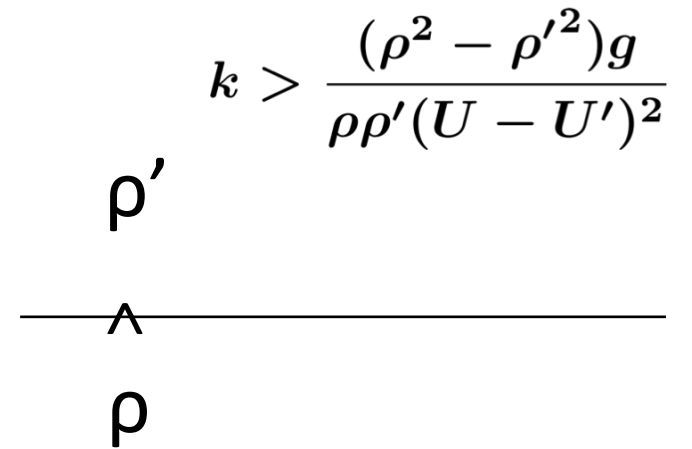
surface gravity wave



Rayleigh-Taylor instability



Kelvin-Helmholtz instability



7.5 Jeans Instability

perturbation

$$\rho : \rho_0 \rightarrow \rho_0 + \rho_1$$

- **Let us analyze instability of a uniform gas bound by self-gravity**

- Poisson eq. for gravity ... (7.54, 57)
- continuity eq. ... (7.55) (using to eliminate **P**)
- Euler eq. + hydrostatic eq. ... (7.56)

let all variables have following x and t dependences:

$$\exp[i (kx - \omega t)]$$

then we obtain dispersion relation ... (7.58, 59)

$$\omega^2 = c_s^2 (k^2 - k_J^2) , \quad k_J^2 = \frac{4\pi G \rho_0}{c_s^2}$$

Jeans Mass

$$M_J = 4/3 \pi \lambda_J^3 \rho_0$$

perturbations will grow if the size of them is larger than $\lambda_J \equiv 2\pi/k_J$!

7.6 Stellar Oscillations. Helioseismology

- **Let us analyze radial pulsation of stars**

from Euler eq., we obtain the acceleration of a fluid element ... (7.61)

$$\frac{dv}{dt} = -\frac{1}{\rho} \frac{dp}{dr} - \frac{GM}{r^2}$$

assuming uniform expansion and considering only linear term,

$$dp/dr = -3 \rho_0/r_0, \quad dp/dr = \gamma (dp/dr) \rho_0/\rho_0 = -3\gamma \rho_0/r_0$$

substituting these to (7.61) and eliminate p by using hydrostatic eq.,
we obtain (7.64)

$$\frac{dv}{dt} = -\frac{GM}{r_0^2} 3\delta \left(\gamma - \frac{4}{3} \right)$$

in the case of $\delta > 0$, when $\gamma > 4/3$ acceleration become negative, so that stars are stable.

